

GRB 060218/SN 2006AJ: A GAMMA-RAY BURST AND PROMPT SUPERNOVA AT  $Z = 0.0335$ N. MIRABAL<sup>1</sup>, J. P. HALPERN<sup>2</sup>, D. AN<sup>3</sup>, J. R. THORSTENSEN<sup>4</sup>, & D. M. TERNDROP<sup>3</sup>*Draft version February 5, 2008*

## ABSTRACT

We report the imaging and spectroscopic localization of GRB 060218 to a low-metallicity dwarf starburst galaxy at  $z = 0.03345 \pm 0.00006$ . In addition to making it the second nearest gamma-ray burst known, optical spectroscopy reveals the earliest detection of weak, supernova-like Si II near 5720 Å ( $\sim 0.1c$ ), starting 1.95 days after the burst trigger. *UBVRI* photometry obtained between 1 and 26 days post-burst confirms the early rise of supernova light, and suggests a short time delay between the gamma-ray burst and the onset of the supernova explosion if the early appearance of a soft component in the X-ray spectrum is understood as a “shock breakout.” Together, these results verify the long-hypothesized origin of soft gamma-ray bursts in the deaths of massive stars.

*Subject headings:* gamma rays: bursts – supernovae: individual (SN 2006aj) – supernovae: general

## 1. INTRODUCTION

It is now accepted that the so-called long-soft ( $\gtrsim 2$  s) gamma-ray bursts (GRBs) accompany some core-collapse supernovae of Type Ic (Galama et al. 1998; Patat et al. 2001; Stanek et al. 2003; Hjorth et al. 2003). The collective evidence also lends credence to the collapsar model for GRBs, in which a relativistic jet breaks through and explodes a hydrogen-stripped Wolf-Rayet star (Woosley 1993; MacFadyen & Woosley 1999). Unfortunately, our understanding of these energetic explosions is still limited by the paucity of nearby ( $z \lesssim 0.2$ ) GRBs with high-quality photometry and spectroscopy.

On UT 2006 February 18.149, the *Swift* Burst Alert Telescope (BAT) detected an unusually long duration high-energy event (Cusumano et al. 2006a). Its prompt gamma-ray light curve was soft, confined to the 15–50 keV band for the first 290 s (Barbier et al. 2006); this was followed by a spectrally harder (25–100 keV) 10-s “spike” that concluded with an exponential coda extending beyond  $t \approx 2000$  s (Barthelmy et al. 2006). The relatively odd behavior of this high-energy transient generated uncertainty as to its basic nature (Gehrels 2006, and references therein), and poor observing conditions at ground-based observatories prevented a quick resolution. However, the identification of an extended optical object at the precise position of the X-ray and optical transient (OT) in pre-burst observations of this field (Cool et al. 2006; Mirabal 2006) favored an extragalactic location. Ultimately, optical spectroscopic observations of the OT discovered by the *Swift* UV/Optical Telescope (Cusumano et al. 2006a) confirmed the low-redshift  $z = 0.033$ , extragalactic nature of this unusual GRB (Mirabal & Halpern 2006a), making it the second lowest burst redshift known to date after GRB 980425/SN 1998bw at  $z = 0.0085$ .

In this Letter we describe the identification of the GRB

060218 host galaxy and its redshift, together with photometry and spectroscopy that verify its origin in the explosion of a massive star, and discuss the implications of our results. We assume an  $H_0 = 71$  km s<sup>−1</sup> Mpc<sup>−1</sup>,  $\Omega_m = 0.27$ ,  $\Omega_\Lambda = 0.73$  cosmological model, corresponding to a luminosity distance  $D_L = 145$  Mpc at  $z = 0.0335$  and an angular scale 0.658 kpc arcsec<sup>−1</sup>.

## 2. OBSERVATIONS

Optical observations of GRB 060218 began at the MDM Observatory on Feb. 19.146 UT using the 2.4m telescope and RETROCAM, the Retractable Optical Camera for Monitoring, equipped with SDSS filters (Morgan et al. 2005), and continued on February 20. Additional *UBVRI* photometric observations were carried out on the MDM 1.3m and 2.4m telescopes using a SITE 2048 × 2048 thinned, back side-illuminated CCD on several nights from February 21 to March 16. All of the photometry was converted to a common *UBVRI* system using Landolt (1992) standard stars, and corrected for Galactic extinction assuming  $E(B - V) = 0.142$  from the dust maps of Schlegel, Finkbeiner, & Davis (1998). We note that this value is consistent with the extinction estimated from high-resolution spectra by Guenther et al. (2006) using the combined Na I D absorption-line equivalent widths from the Galaxy and GRB host. This uniform data set, listed in Table 1 and shown in Fig. 1, can be fitted with a power-law decay plus a supernova (SN) light curve that will be described in more detail in the following section.

The position of the OT in an MDM 2.4m image was measured with respect to the apparent host galaxy SDSS J032139.68+165201.7 using a set of unsaturated stars common to both images. We find that the OT is centered on the compact galaxy to less than 0".2 (130 pc) in each coordinate. This is to be compared with the galaxy’s half-light radius,  $r_{1/2} \approx 1".5$  (1.0 kpc).

Spectra of GRB 060218 were obtained starting on February 20.097 UT with the Boller & Chivens CCD spectrograph (CCDS) mounted on the 2.4 m telescope. The setup used provides 3.1 Å pixel<sup>−1</sup> dispersion and  $\approx 8.2$  Å resolution with a 1" slit. The observations consisted of two 1800 s integrations under fair sky conditions. The spectra were processed using standard procedures in

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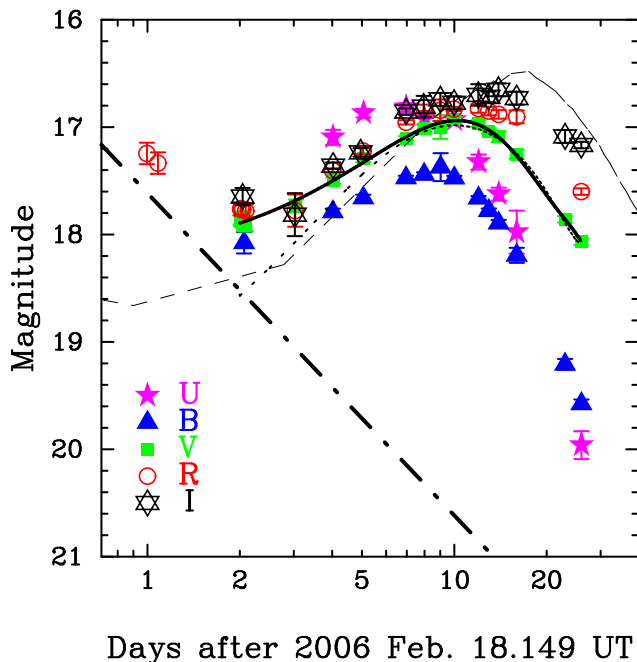


FIG. 1.— *UBVRI* data for GRB 060218, corrected for Galactic extinction and host-galaxy contamination. The *solid* line is a fit to the *V*-band light curve. The *dotted* line is a fit to the *V*-band light curve after subtracting an  $\alpha = 1.2$  power-law decay (*dot-dashed* line) as justified in the text. The *dashed* line is a template of the *V*-band light curve of SN 1998bw (Galama et al. 1998) shifted to  $z = 0.0335$ . [See the electronic edition of the *Journal* for a color version of this figure.]

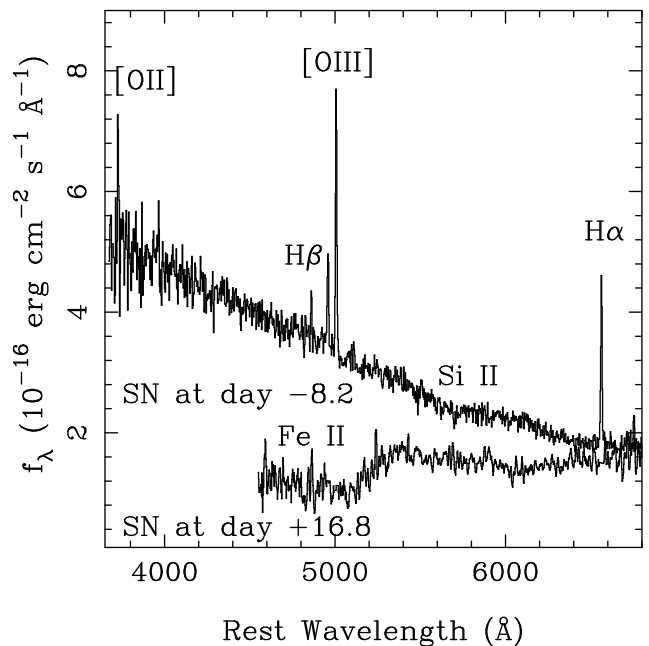


FIG. 2.— Spectra of GRB 060218 obtained on 2006 February 20.097 UT (1.95 days after the burst) and March 17.12 (27 days post-burst). Days relative to supernova maximum are indicated. Starburst emission lines from the host galaxy were excised from the second spectrum. The first spectrum marks the earliest appearance of Si II near 5720 Å, while its continuum is reasonably well fitted by a spectral index  $\beta = 0.1 \pm 0.3$  (see text). Photometry before and after the spectrum was taken (Table 1) is consistent with this spectral index.

IRAF.<sup>5</sup> The wavelength scale was established by fitting a set of polynomials to Xe lamp spectra obtained immediately after each target exposure. The spectrophotometric standard star Feige 34 (Stone 1977), observed at comparable telescope pointing to the GRB, was used for flux calibration. Although no order-separating filter was used, we expect that second-order contamination is less than 1.5% below 7000 Å (e.g., Izotov et al. 2001). Another set of spectra, consisting of three 720 s integrations, was obtained on March 17.12 UT using the Modular Spectrograph (ModSpec) on the 2.4 m telescope, which provides 2 Å pixel<sup>-1</sup> dispersion and  $\approx 3.6$  Å resolution with a 1''/1 slit. Similar reduction steps, plus correction for atmospheric absorption bands, were performed. Figure 2 shows the dereddened (Cardelli, Clayton, & Mathis 1989), wavelength- and flux-calibrated spectra of GRB 060218/SN 2006aj.

### 3. RESULTS

Strong, redshifted nebular emission lines identified in the spectrum are listed in Table 2. The fluxes are taken from the CCDS spectrum, while the more accurate wavelengths from the ModSpec are listed. The weighted mean heliocentric redshift derived from the emission lines is  $z = 0.03345 \pm 0.00006$ , and the line widths are unresolved at the resolution of  $\approx 160$  km s<sup>-1</sup>. The emission-line redshift is consistent with the Na I D absorption-line

velocities from the host (Guenther et al. 2006), falling within the 20 km s<sup>-1</sup> spread of the latter. The Balmer decrement indicates little or no intrinsic reddening. In addition, the earlier spectrum reveals a relatively blue continuum that is reasonably well fitted by a power-law of the form  $f_\nu \propto \nu^{-\beta}$  with spectral index  $\beta = 0.1 \pm 0.3$  and by a broad P Cygni feature with the bottom of the absorption trough at  $\approx 5720$  Å rest wavelength.

We argue that this spectral feature corresponds to Si II  $\lambda 6355$  with an expansion velocity  $31500 \pm 9200$  km s<sup>-1</sup>. An identification with Na I D moving at about 8700 km s<sup>-1</sup>, or with He I  $\lambda 5876$  at even lower velocity, appears less likely when compared to the early-time optical spectra of Type Ic SNe (Patat et al. 2001). This detection signals the emergence of the supernova designated SN 2006aj (Masetti et al. 2006; Soderberg, Berger, & Schmidt 2006; Fugazza et al. 2006; Mirabal & Halpern 2006b; Fatkhullin et al. 2006; Mazzali & Pian 2006; Modjaz et al. 2006), and seals the connection between GRB 060218 and the explosion of a massive star. The weak Si II line, as well as a nearly featureless continuum at early times, are typical of Type Ic SN explosions lacking both a hydrogen and helium envelope (e.g. Filippenko 1997). SN 2006aj appears to be much bluer than the Type Ic SN 1998bw. The early spectral shape of SN 2006aj resembles more closely the Type Ic GRB 030329/SN 2003dh spectra obtained within a week of that burst (Stanek et al. 2003; Hjorth et al. 2003).

The emergence of a weak, broad Si II feature so soon after the *Swift* BAT localization implies that SN 2006aj

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began  $< 1.95$  days after the GRB 060218. This basic picture is confirmed by the flattening and rise in the optical light curve on 2006 February 21.181 UT (Fig. 1). To better constrain the SN contribution to the light curve, we subtracted the host galaxy contribution and, optionally, a power-law decay assumed to be from relativistic ejecta interacting with the circumburst medium. The host galaxy magnitudes were transformed from Cool et al. (2006), Adelman-McCarthy et al. (2006), and our own Landolt star calibrations. There is direct evidence that emission generated by the relativistic ejecta began on the first day after the trigger, when a temporal power-law index  $\alpha \approx 1.2$  describes the X-ray observations (Cusumano et al. 2006b). Such decay is slightly steeper than a power-law index  $\alpha \approx 0.7 \pm 0.3$  that can be fitted to our  $R$ -band measurements prior to February 21, but these are consistent because the optical band is already affected by the SN rise at this point.

Taking  $\alpha \approx 1.2$  as an estimate for the early optical decay rate (the *dot-dashed line* in Fig. 1), and subtracting it and the contribution of the host galaxy will produce the residual  $V$ -band SN light curve shown as the *dotted line* in Figure 1. Allowing for uncertainties in the initial optical decay rate, which could be slower than in the X-ray, the SN rise is consistent with an origin at the GRB trigger time. Alternatively, if we ignore the  $R$ -band measurements obtained on February 19–20, and assume that the light curve is completely dominated by the SN 2006aj rise after February 21, we get a much flatter SN light curve (the solid line in Fig. 1). The latter model is less attractive, as it neglects the first 2 days of bright, decaying optical emission and points to a supernova time several days before the GRB.

For comparison we show a fit of the  $V$ -band light curve of SN 1998bw (Galama et al. 1998) shifted to  $z = 0.0335$ . Both the raw and modeled  $V$ -band light curves suggest that SN 2006aj reached maximum in  $V$  on 2006 February 28.3 UT ( $10.2 \pm 0.3$  days after the burst). The peak of SN 2006aj occurs earlier than in SN 1998bw, and is more like the Type Ic SN 2002ap (Gal-Yam, Ofek, & Shemmer 2002; Mazzali et al. 2002). The peak absolute magnitude of SN 2006aj is  $M_V = -18.87$ . Although it was much bluer than SN 1998bw early on, SN 2006aj’s maximum is 0.53 mag fainter than SN 1998bw, which probably translates into  $\lesssim 0.5 M_\odot$  of ejected  $^{56}\text{Ni}$  mass during the explosion (Iwamoto et al. 1998; Woosley, Eastman, & Schmidt 1999).

At a redshift of  $z = 0.0335$  the  $\gamma$ -ray fluence of GRB 060218 corresponds to isotropic energy  $E_{\text{iso}} = (6.2 \pm 0.3) \times 10^{49}$  ergs, and the peak luminosity is  $L_p = (5 \pm 3) \times 10^{46}$  ergs  $\text{s}^{-1}$  (Sakamoto et al. 2006; Campana et al. 2006). This energy release is at least an order of magnitude lower than the average energy measured in long-duration GRBs and yet a factor of  $\sim 20$  larger than the intrinsically weak GRB 980425/SN 1998bw (Galama et al. 1998). This energy is in fact comparable to the soft X-ray flash XRF 020903 (Sakamoto et al. 2006). In terms of the X-ray emission, the isotropic X-ray luminosity of GRB 060218 at  $t = 10$  hr is  $L_X \sim 10^{43}$  ergs  $\text{s}^{-1}$  (Cusumano et al. 2006b). This is a factor of  $10^3$  fainter than the sample of GRB X-ray luminosities culled by Berger, Kulkarni, & Frail (2003). However, this strict comparison may not be meaningful, since it does not take

into account the contribution from an earlier flaring period. In fact, a rough estimate of the X-ray fluence of GRB 060218 for the first orbit of *Swift* data (159–2770 s post-trigger) yields  $\approx 2.3 \times 10^{49}$  ergs, a large fraction of the total.

Next, we examine the host galaxy of GRB 060218 in greater detail. The observed line ratios are typical of a high-excitation starburst galaxy. In particular, the measured  $\text{H}\alpha$  line flux, uncorrected for extinction at the host galaxy, implies a star-formation rate (SFR) equivalent to  $\approx 0.05\text{--}0.15 M_\odot \text{ yr}^{-1}$  (Kennicutt 1998). A similar computation of the SFR using  $[\text{O II}]$  yields  $0.09 \pm 0.05 M_\odot \text{ yr}^{-1}$ . We also find a good agreement with the SFR derived using an extrapolation of the UV continuum luminosity of the host galaxy as tabulated in Cool et al. (2006).

Turning our attention to host galaxy metallicity, we can determine relative abundances from emission-line strengths of  $[\text{O II}]$ ,  $\text{H}\beta$ , and  $[\text{O III}]$  (Table 2) by adopting the calibrations in Kobulnicky & Kewley (2004). Under this approximation, the metallicity is  $[\text{O/H}] = -0.34 \pm 0.3$ , assuming  $\log(\text{O/H})_\odot + 12 = 8.72$  (Allende Prieto et al. 2001). This value is slightly larger than the measurement for the XRF 031203 host (Prochaska et al. 2004), but it is still among the lowest observed for GRB hosts. We also estimate  $M_B \approx -16.01$  for SDSS J032139.68+165201.7, which ranks this galaxy in the bright end of Local Group dwarf galaxies (Mateo 1998).

#### 4. DISCUSSION

Together, its energetics and host-galaxy metallicity place GRB 060218 somewhere between GRB 980425/SN 1998bw (Galama et al. 1998) and XRF 031203 (Soderberg et al. 2004). Most importantly, this event continues to shape a picture in which sub-luminous, sub-energetic GRB-supernovae born in low-metallicity, dwarf galaxies dominate the local ( $z \lesssim 0.5$ ) population of GRB events (e.g. Soderberg et al. 2004). It remains to be seen whether this empirical result constitutes a selection effect (i.e., we are simply missing the analogues of faint explosions at higher redshifts) or whether the trend is indeed a real consequence of stellar and/or metallicity evolution over the ages (Ramirez-Ruiz, Lazzati, & Blain 2002; Woosley & Janka 2005; Langer & Norman 2006). If the latter is correct, low-metallicity progenitors in the local Universe may differ from GRB progenitors at higher redshifts.

It is notable that during the early, X-ray bright phase the absorbing column density needed to fit the *Swift* X-ray spectrum,  $N_{\text{H}} = 6 \times 10^{21} \text{ cm}^{-2}$  (Campana et al. 2006), is considerably greater than both the Galactic and host extinction derived from optical emission and absorption lines, as well as from the optical colors of the afterglow. The standard Galactic ratio  $N(\text{H I} + \text{H}_2)/E(B - V) = 5.8 \times 10^{21} \text{ cm}^{-2} \text{ mag}^{-1}$  (Bohlin, Savage, & Drake 1978) would predict  $E(B - V) \approx 1$ , rather than the  $E(B - V) \approx 0.17 \pm 0.03$  observed from the optical methods. Since “X-ray  $N_{\text{H}}$ ” is not really  $N_{\text{H}}$ , but a proxy for the heavier elements that dominate X-ray photoelectric absorption, this implies a dust-deficient medium. The stellar wind of a Wolf-Rayet progenitor has enough column density to be the location of this excess photoelectric X-ray absorption.

Phenomenologically, GRB 060218 does share some properties with the “classical” high-redshift burst population, showing the canonical imprint of emission from a relativistic blast wave or jet running into circumstellar material ( $\alpha \approx 1.2$ ) prior to the SN emergence. It also reveals extended prompt X-ray emission remarkably similar to the early X-ray light curves inferred from *Swift* X-ray Telescope observations (Zhang et al. 2005). Based solely on the early results from GRB 060218, and using our sparse optical data before the SN rise, we cannot firmly distinguish among a relativistic jet afterglow, emission from a jet cocoon (Ramirez-Ruiz, Celotti, & Rees 2002), or decaying blackbody radiation associated with “shock breakout” (Campana et al. 2006). As calculated in spherical SN models, shock breakout refers to the first observable event after core collapse, and it will occur at  $\Delta t \sim 1$  hr (Arnett 1996). But in the likely case of a highly asymmetric explosion accompanying a GRB, a narrow “jet breakout” can be observed, as well as a jet-driven shock emerging with different delays from around the stellar envelope.

## 5. CONCLUSIONS

We reported the identification of GRB 060218 as the second nearest GRB known to date. Taken together, the emerging SN light curve, the development of a Si II feature, and the lack of hydrogen lines in the spectrum indicate that the progenitor of this event was a massive star, most likely in the Wolf-Rayet family, that was stripped of its hydrogen and helium envelope prior to the explosion. The presence of early ( $\lesssim 2$  days after the burst) residual light in the decay suggests that a fraction of the early emission was created by a mildly relativistic blast wave or jet running into circumstellar material. Furthermore, the extrapolation of a derived SN 2006aj light curve back in time supports the idea that the GRB was nearly simultaneous with the massive core collapse that gave rise to the SN. Finally, the sub-energetic nature of this nearby event and its location within a low-metallicity host galaxy highlight the possibility that the variety in massive stellar explosions is, in part, intrinsic to the metallicity of their progenitors.

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TABLE 1  
OPTICAL PHOTOMETRY OF GRB 060218

Date (UT)	Filter	Magnitude <sup>a</sup>	Date (UT)	Filter	Magnitude <sup>a</sup>	Date (UT)	Filter	Magnitude <sup>a</sup>
2006 Feb. 22.173	<i>U</i>	17.10 ± 0.07	2006 Feb. 20.208	<i>V</i>	17.89 ± 0.06	2006 Feb. 25.114	<i>R</i>	16.96 ± 0.02
2006 Feb. 23.210	<i>U</i>	16.87 ± 0.04	2006 Feb. 21.196	<i>V</i>	17.73 ± 0.06	2006 Feb. 26.112	<i>R</i>	16.82 ± 0.03
2006 Feb. 25.119	<i>U</i>	16.82 ± 0.04	2006 Feb. 22.186	<i>V</i>	17.50 ± 0.03	2006 Feb. 27.162	<i>R</i>	16.84 ± 0.10
2006 Feb. 26.154	<i>U</i>	16.92 ± 0.05	2006 Feb. 23.196	<i>V</i>	17.30 ± 0.03	2006 Feb. 28.136	<i>R</i>	16.83 ± 0.02
2006 Feb. 28.160	<i>U</i>	16.93 ± 0.04	2006 Feb. 25.112	<i>V</i>	17.10 ± 0.02	2006 Mar. 2.140	<i>R</i>	16.83 ± 0.04
2006 Mar. 2.153	<i>U</i>	17.32 ± 0.07	2006 Feb. 26.140	<i>V</i>	17.01 ± 0.03	2006 Mar. 3.098	<i>R</i>	16.85 ± 0.04
2006 Mar. 4.106	<i>U</i>	17.62 ± 0.06	2006 Feb. 27.167	<i>V</i>	17.00 ± 0.11	2006 Mar. 4.094	<i>R</i>	16.88 ± 0.04
2006 Mar. 6.135	<i>U</i>	17.98 ± 0.20	2006 Feb. 28.168	<i>V</i>	16.93 ± 0.02	2006 Mar. 6.113	<i>R</i>	16.90 ± 0.06
2006 Mar. 16.111	<i>U</i>	19.96 ± 0.13	2006 Mar. 2.143	<i>V</i>	16.97 ± 0.03	2006 Mar. 16.106	<i>R</i>	17.60 ± 0.04
2006 Feb. 20.214	<i>B</i>	18.08 ± 0.10	2006 Mar. 3.102	<i>V</i>	17.04 ± 0.02	2006 Feb. 20.191	<i>I</i>	17.65 ± 0.08
2006 Feb. 22.169	<i>B</i>	17.79 ± 0.03	2006 Mar. 4.097	<i>V</i>	17.09 ± 0.05	2006 Feb. 21.172	<i>I</i>	17.82 ± 0.20
2006 Feb. 23.202	<i>B</i>	17.66 ± 0.03	2006 Mar. 6.118	<i>V</i>	17.25 ± 0.04	2006 Feb. 22.177	<i>I</i>	17.36 ± 0.03
2006 Feb. 25.109	<i>B</i>	17.47 ± 0.02	2006 Mar. 13.112	<i>V</i>	17.86 ± 0.03	2006 Feb. 23.100	<i>I</i>	17.24 ± 0.04
2006 Feb. 26.126	<i>B</i>	17.44 ± 0.02	2006 Mar. 16.108	<i>V</i>	18.07 ± 0.03	2006 Feb. 25.116	<i>I</i>	16.86 ± 0.03
2006 Feb. 27.173	<i>B</i>	17.37 ± 0.13	2006 Feb. 19.146	<i>R</i>	17.25 ± 0.10	2006 Feb. 26.102	<i>I</i>	16.81 ± 0.10
2006 Feb. 28.015	<i>B</i>	17.47 ± 0.02	2006 Feb. 19.230	<i>R</i>	17.34 ± 0.10	2006 Feb. 27.158	<i>I</i>	16.74 ± 0.06
2006 Mar. 2.147	<i>B</i>	17.66 ± 0.02	2006 Feb. 20.162	<i>R</i>	17.76 ± 0.06	2006 Feb. 28.134	<i>I</i>	16.77 ± 0.06
2006 Mar. 3.105	<i>B</i>	17.78 ± 0.02	2006 Feb. 20.168	<i>R</i>	17.76 ± 0.06	2006 Mar. 2.130	<i>I</i>	16.70 ± 0.10
2006 Mar. 4.100	<i>B</i>	17.89 ± 0.03	2006 Feb. 20.191	<i>R</i>	17.77 ± 0.06	2006 Mar. 3.089	<i>I</i>	16.72 ± 0.05
2006 Mar. 6.124	<i>B</i>	18.19 ± 0.07	2006 Feb. 20.240	<i>R</i>	17.78 ± 0.06	2006 Mar. 4.091	<i>I</i>	16.65 ± 0.06
2006 Mar. 13.114	<i>B</i>	19.20 ± 0.05	2006 Feb. 21.180	<i>R</i>	17.78 ± 0.15	2006 Mar. 6.109	<i>I</i>	16.73 ± 0.06
2006 Mar. 16.109	<i>B</i>	19.58 ± 0.04	2006 Feb. 22.161	<i>R</i>	17.40 ± 0.02	2006 Mar. 13.109	<i>I</i>	17.09 ± 0.06
2006 Feb. 20.162	<i>V</i>	17.84 ± 0.06	2006 Feb. 23.192	<i>R</i>	17.22 ± 0.03	2006 Mar. 16.105	<i>I</i>	17.17 ± 0.04

<sup>a</sup>Host galaxy with  $U = 20.10$ ,  $B = 20.41$ ,  $V = 20.09$ ,  $R = 19.91$ ,  $I = 19.54$  was subtracted, then remainder was corrected for Galactic extinction  $A_U = 0.77$ ,  $A_B = 0.61$ ,  $A_V = 0.47$ ,  $A_R = 0.38$ , and  $A_I = 0.28$  respectively.

TABLE 2  
HOST-GALAXY EMISSION LINES

Line $\lambda_{\text{air}}(\text{\AA})$	$\lambda_{\text{helio}}(\text{\AA})$	$F(\lambda)/F(\text{H}\beta)^{\text{a}}$
[O II] 3727.5	(3850.7) <sup>b</sup>	2.61 ± 0.4
H $\beta$ 4861.33	5024.36	1.00 ± 0.3
[O III] 4958.92	5124.84	2.03 ± 0.3
[O III] 5006.85	5174.31	5.09 ± 0.6
H $\alpha$ 6562.79	6782.25	2.94 ± 0.3
[N II] 6583.39	...	< 0.08
[S II] 6716.42	...	< 0.65
[S II] 6730.78	...	< 0.54

<sup>a</sup>Flux relative to  $F(\text{H}\beta) = 9.94 \times 10^{-16} \text{ ergs cm}^{-2} \text{ s}^{-1}$ , corrected for Galactic extinction  $E(B - V) = 0.142$ .

<sup>b</sup>Poor wavelength calibration in this region.